

# On the connection of major mergers and AGNs

A closer look at high-accretion rate AGNs at  $z \sim 2$

SMBH: Environment and Evolution • Corfu • 20.06.2019

Victor Marian • MPIA

Advisor: Knud Jahnke

---

# How to trigger and feed a black hole?

# How to trigger and feed a black hole?

- Fuel → Gas
- Transport → Torques

# How to trigger and feed a black hole?

- Fuel → Gas
- Transport → Torques

→ Major mergers **most feasible** option (?)

# Recent studies

show **NO** enhancement in merger incidents for

- the majority of X-ray selected and optical observed AGNs across cosmic time ( $z \leq 1$ ) (Gabor et al. 2009; Georgakakis et al. 2009; Cisternas et al. 2011)
- AGNs at low or intermediate luminosities ( $L_X \leq 10^{43} \text{erg s}^{-1}$ ) (Grogin et al. 2005; Allevato et al. 2011; Schawinski et al. 2011; Kocevski et al. 2012; Böhm et al. 2013; Cheung et al. 2015; Cisternas et al. 2015; Rosario et al. 2015; Goulding et al. 2017)
- AGNs with high luminosities ( $L_X \geq 10^{43} \text{erg s}^{-1}$ ) (Karouzos et al. 2014; Villforth et al. 2014, 2017; Hewlett et al. 2017)
- Black holes with the highest masses (Mehtley et al. 2016)
- Heavily obscured AGNs (Schawinski et al. 2012)
- AGNs in early evolutionary stages (Villforth et al. 2018)

# Recent studies

show **NO** enhancement in merger incidents for

- the majority of X-ray selected and optical observed AGNs across cosmic time ( $z \leq 1$ ) (Gabor et al. 2009; Georgakakis et al. 2009; Cisternas et al. 2011)
- AGNs at low or intermediate luminosities ( $L_X \leq 10^{43} \text{erg s}^{-1}$ ) (Grogin et al. 2005; Alleinato et al. 2011; Schawinski et al. 2011; Kocevski et al. 2012; Böhm et al. 2013; Cheung et al. 2015; Cisternas et al. 2015; Rosario et al. 2015; Goulding et al. 2017)
- AGNs with high luminosities ( $L_X \geq 10^{43} \text{erg s}^{-1}$ ) (Karouzos et al. 2014; Villforth et al. 2014, 2017; Hewlett et al. 2017)
- Black holes with the highest masses (Mehtley et al. 2016)
- Heavily obscured AGNs (Schawinski et al. 2012)
- AGNs in early evolutionary stages (Villforth et al. 2018)

# Recent studies

show **NO** enhancement in merger incidents for

- the majority of X-ray selected and optical observed AGNs across cosmic time ( $z \leq 1$ ) (Gabor et al. 2009; Georgakakis et al. 2009; Cisternas et al. 2011)
- AGNs at low or intermediate luminosities ( $L_x \leq 10^{43} \text{erg s}^{-1}$ ) (Grogin et al. 2005; Alleinato et al. 2011; Schawinski et al. 2011; Kocevski et al. 2012; Böhm et al. 2013; Cheung et al. 2015; Cisternas et al. 2015; Rosario et al. 2015; Goulding et al. 2017)
- AGNs with high luminosities ( $L_x \geq 10^{43} \text{erg s}^{-1}$ ) (Karouzos et al. 2014; Villforth et al. 2014, 2017; Hewlett et al. 2017)
- Black holes with the highest masses (Mehtley et al. 2016)
- Heavily obscured AGNs (Schawinski et al. 2012)
- AGNs in early evolutionary stages (Villforth et al. 2018)

# Recent studies

show **NO** enhancement in merger incidents for

- the majority of X-ray selected and optical observed AGNs across cosmic time ( $z \leq 1$ ) (Gabor et al. 2009; Georgakakis et al. 2009; Cisternas et al. 2011)
- AGNs at low or intermediate luminosities ( $L_x \leq 10^{43} \text{erg s}^{-1}$ ) (Grogin et al. 2005; Alleinato et al. 2011; Schawinski et al. 2011; Kocevski et al. 2012; Böhm et al. 2013; Cheung et al. 2015; Cisternas et al. 2015; Rosario et al. 2015; Goulding et al. 2017)
- AGNs with high luminosities ( $L_x \geq 10^{43} \text{erg s}^{-1}$ ) (Karouzos et al. 2014; Villforth et al. 2014, 2017; Hewlett et al. 2017)
- Black holes with the highest masses (Mehtley et al. 2016)
- Heavily obscured AGNs (Schawinski et al. 2012)
- AGNs in early evolutionary stages (Villforth et al. 2018)



# Recent studies

show **NO** enhancement in merger incidents for

- the majority of X-ray selected and optical observed AGNs across cosmic time ( $z \leq 1$ ) (Gabor et al. 2009; Georgakakis et al. 2009; Cisternas et al. 2011)
- AGNs at low or intermediate luminosities ( $L_x \leq 10^{43} \text{erg s}^{-1}$ ) (Grogin et al. 2005; Alleinato et al. 2011; Schawinski et al. 2011; Kocevski et al. 2012; Böhm et al. 2013; Cheung et al. 2015; Cisternas et al. 2015; Rosario et al. 2015; Goulding et al. 2017)
- AGNs with high luminosities ( $L_x \geq 10^{43} \text{erg s}^{-1}$ ) (Karouzos et al. 2014; Villforth et al. 2014, 2017; Hewlett et al. 2017)
- Black holes with the highest masses (Mehtley et al. 2016)
- Heavily obscured AGNs (Schawinski et al. 2012)
- AGNs in early evolutionary stages (Villforth et al. 2018)

# Recent studies

show **NO** enhancement in merger incidents for

- the majority of X-ray selected and optical observed AGNs across cosmic time ( $z \leq 1$ ) (Gabor et al. 2009; Georgakakis et al. 2009; Cisternas et al. 2011)
- AGNs at low or intermediate luminosities ( $L_x \leq 10^{43} \text{erg s}^{-1}$ ) (Grogin et al. 2005; Allevato et al. 2011; Schawinski et al. 2011; Kocevski et al. 2012; Böhm et al. 2013; Cheung et al. 2015; Cisternas et al. 2015; Rosario et al. 2015; Goulding et al. 2017)
- AGNs with high luminosities ( $L_x \geq 10^{43} \text{erg s}^{-1}$ ) (Karouzos et al. 2014; Villforth et al. 2014, 2017; Hewlett et al. 2017)
- Black holes with the highest masses (Mehtley et al. 2016)
- Heavily obscured AGNs (Schawinski et al. 2012)
- AGNs in early evolutionary stages (Villforth et al. 2018)

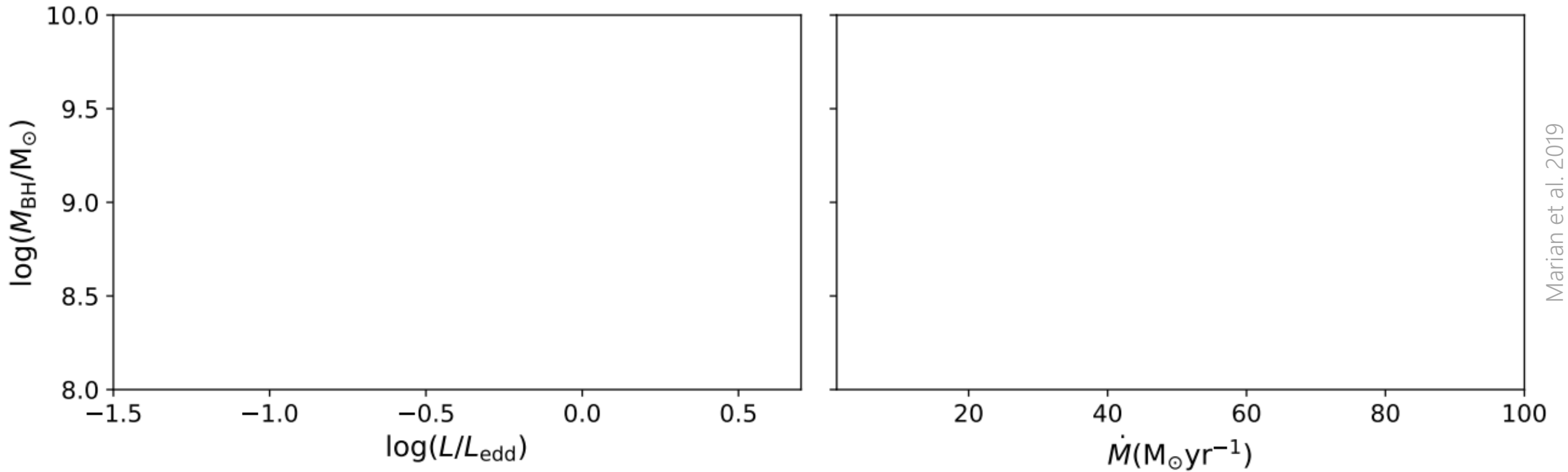
# Recent studies

show **NO** enhancement in merger incidents for

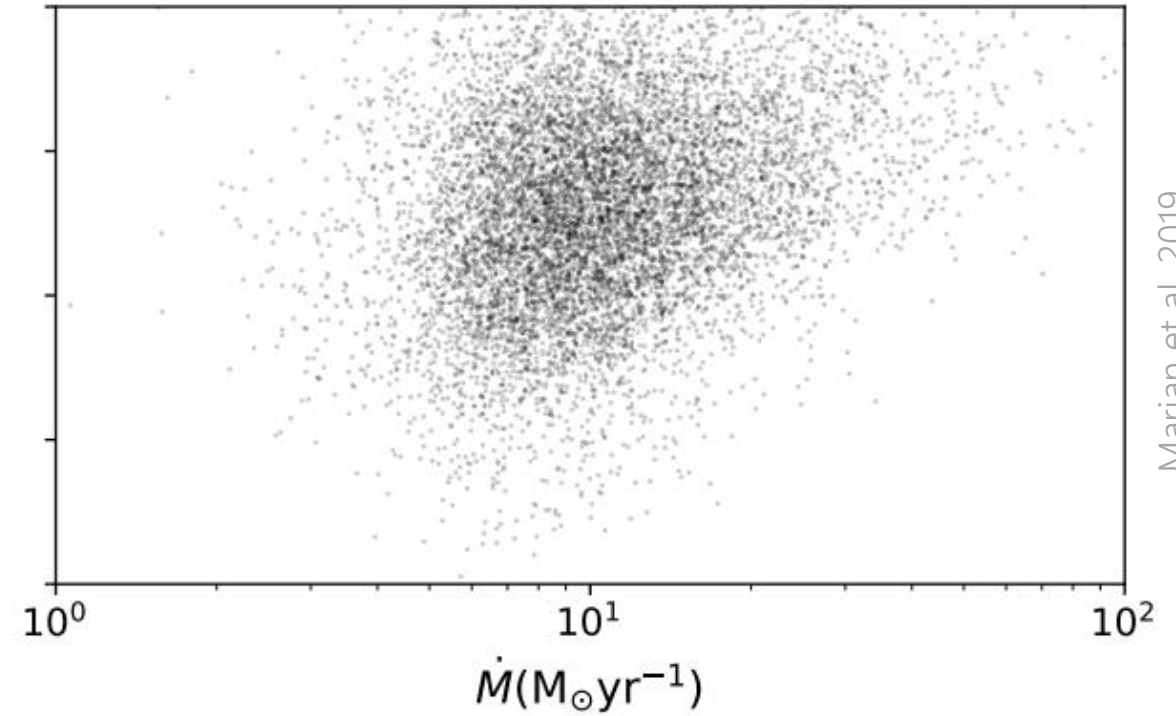
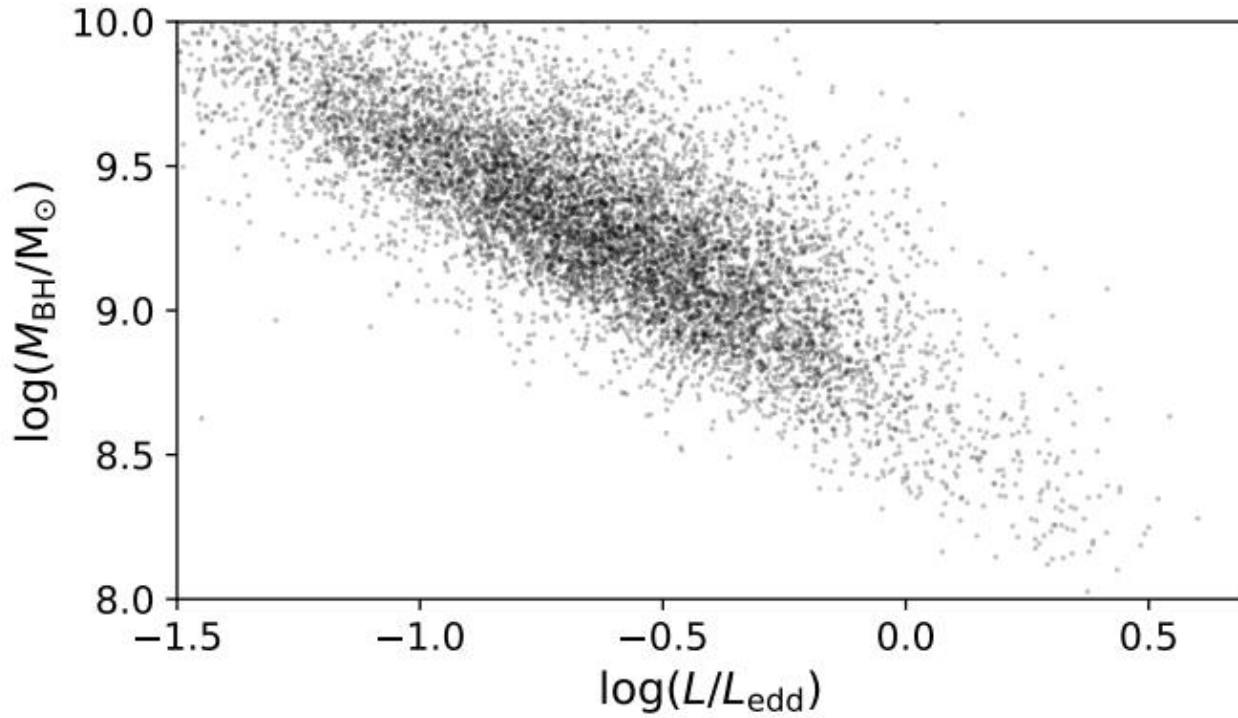
- the majority of X-ray selected and optical observed AGNs across cosmic time ( $z \leq 1$ ) (Gabor et al. 2009; Georgakakis et al. 2009; Cisternas et al. 2011)
- AGNs at low or intermediate luminosities ( $L_x \leq 10^{43} \text{erg s}^{-1}$ ) (Grogin et al. 2005; Allevato et al. 2011; Schawinski et al. 2011; Kocevski et al. 2012; Böhm et al. 2013; Cheung et al. 2015; Cisternas et al. 2015; Rosario et al. 2015; Goulding et al. 2017)
- AGNs with high luminosities ( $L_x \geq 10^{43} \text{erg s}^{-1}$ ) (Karouzos et al. 2014; Villforth et al. 2014, 2017; Hewlett et al. 2017)
- Black holes with the highest masses (Mehtley et al. 2016)
- Heavily obscured AGNs (Schawinski et al. 2012)
- AGNs in early evolutionary stages (Villforth et al. 2018)

# Matching control sample of inactive galaxies essential

# Sample selection



# Sample selection

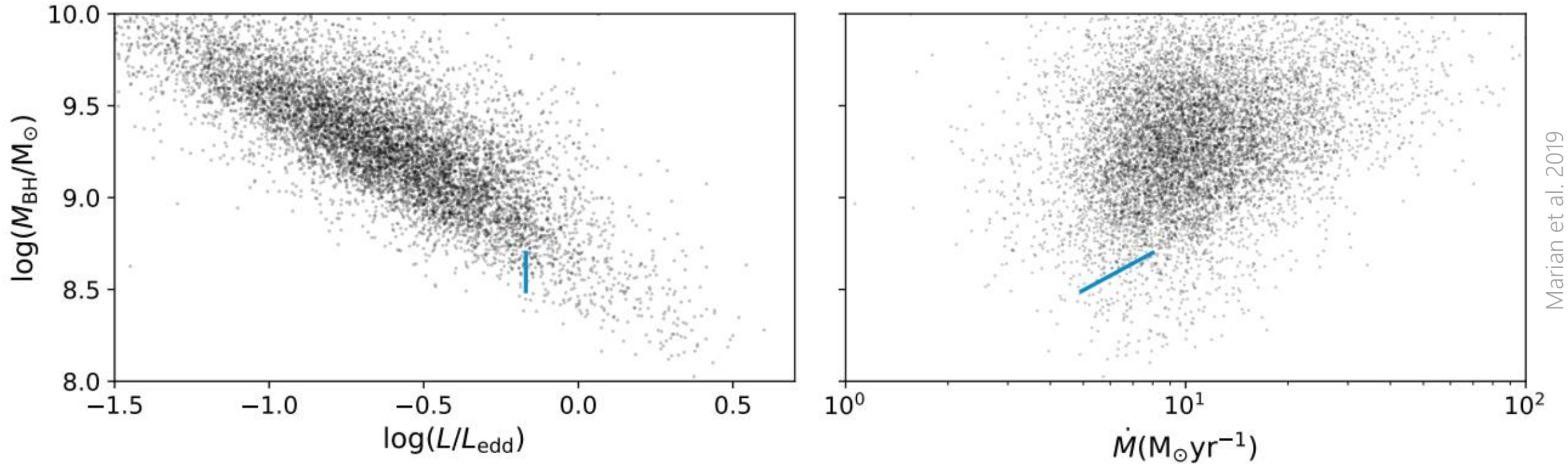


Marian et al. 2019

**Redshift**

$$1.8 \leq z \leq 2.2$$

# Sample selection



Marian et al. 2019

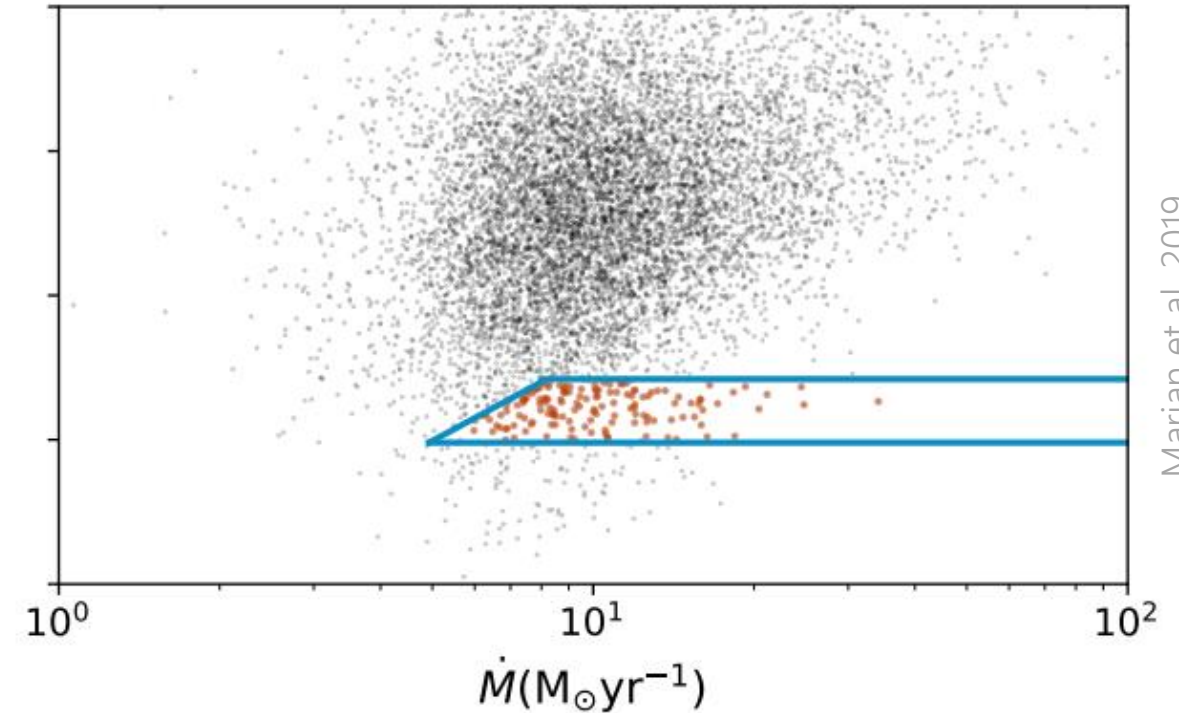
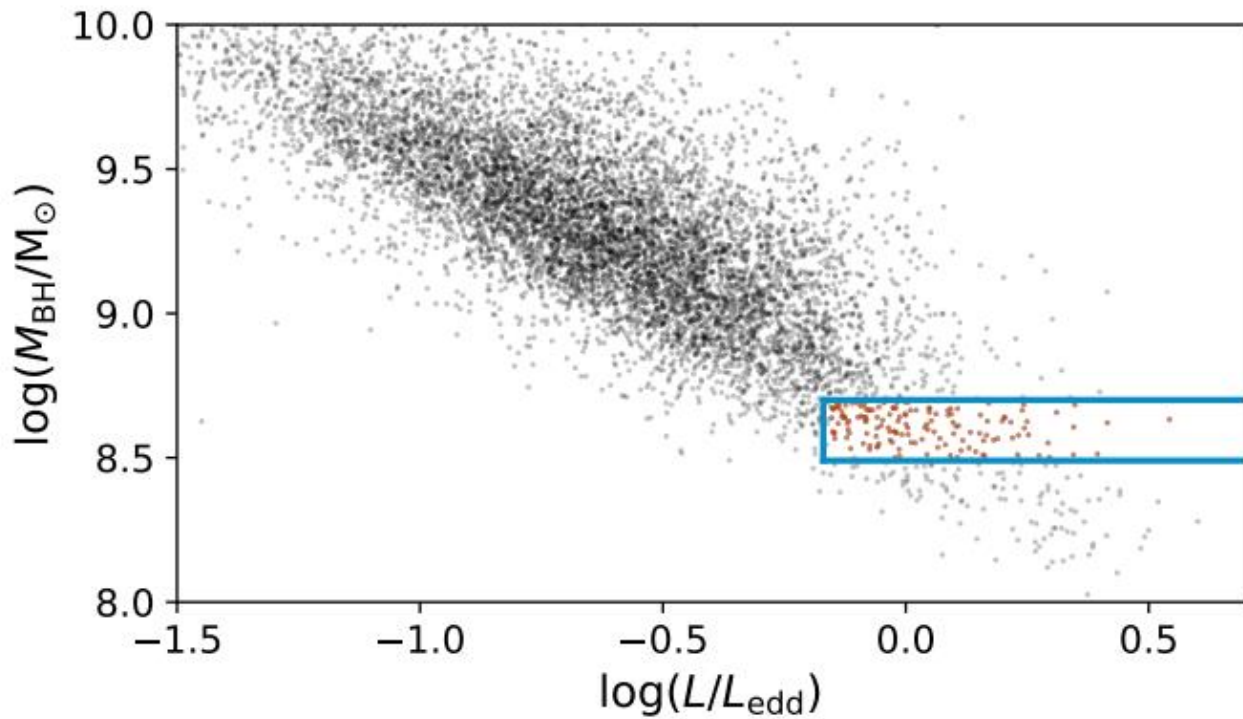
**Redshift**

$$1.8 \leq z \leq 2.2$$

**Mass**

$$8.5 \leq \log(M_{\text{BH}}/M_{\odot}) \leq 8.7$$

# Sample selection



Marian et al. 2019

**Redshift**

$$1.8 \leq z \leq 2.2$$

**Mass**

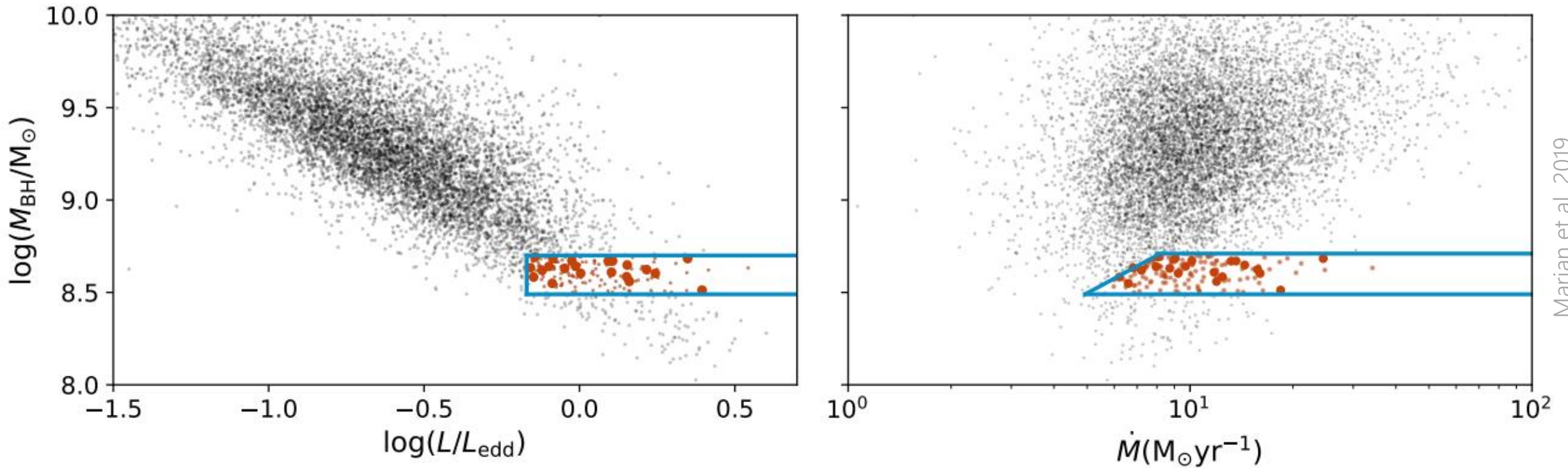
$$8.5 \leq \log(M_{\text{BH}}/M_{\odot}) \leq 8.7$$

**Eddington ratio**

$$(L/L_{\text{edd}}) \geq 70\%$$



# Sample selection



Marian et al. 2019

## Redshift

$$1.8 \leq z \leq 2.2$$

## Mass

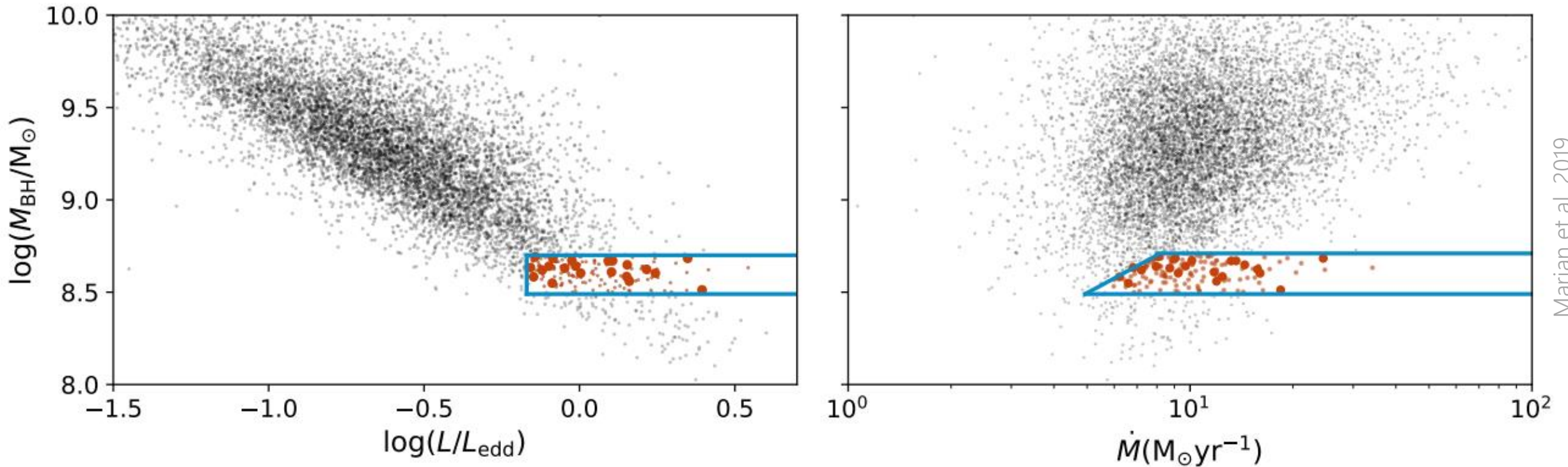
$$8.5 \leq \log(M_{\text{BH}}/M_{\odot}) \leq 8.7$$

Sample: **21 AGN, Type I**

## Eddington ratio

$$(L/L_{\text{edd}}) \geq 70\%$$

# Sample selection



Marian et al. 2019

**Redshift**

$$1.8 \leq z \leq 2.2$$

**Mass**

$$8.5 \leq \log(M_{\text{BH}}/M_{\odot}) \leq 8.7$$

**Eddington ratio**

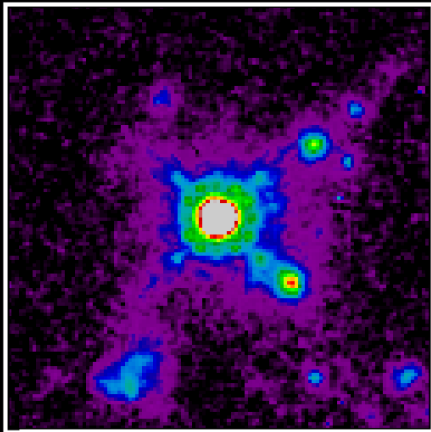
$$(L/L_{\text{edd}}) \geq 70\%$$

Sample: **21 AGN, Type I**

Comparison sample: **92 inactive galaxies** from CANDELS

# Modeling & Point-source subtraction

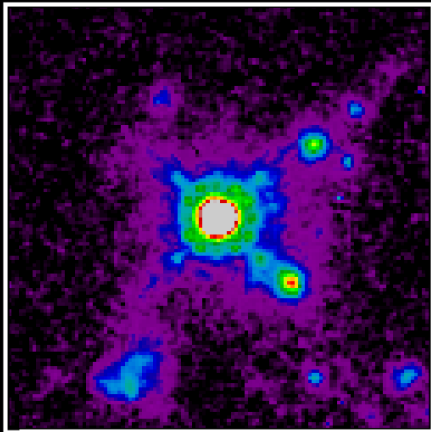
Marian et al. 2019



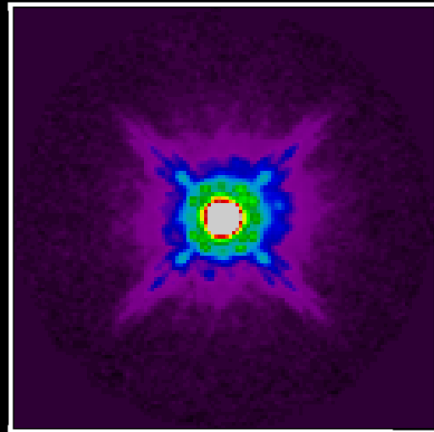
AGN + host  
galaxy

# Modeling & Point-source subtraction

Marian et al. 2019



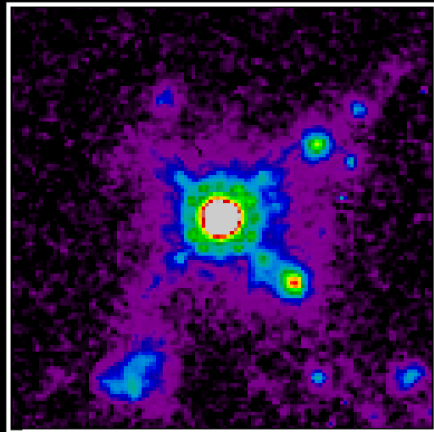
AGN + host  
galaxy



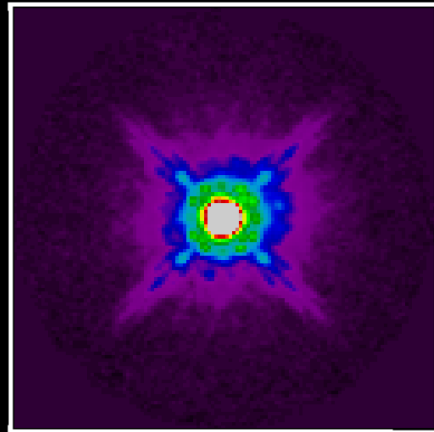
Convolved  
model

# Modeling & Point-source subtraction

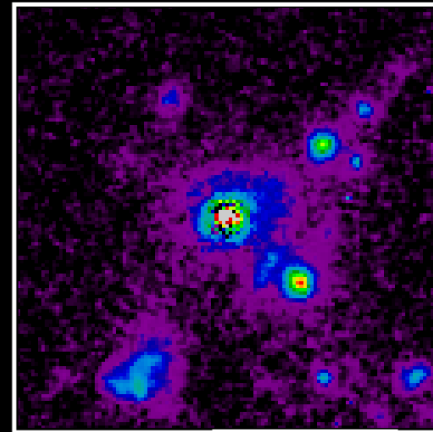
Marian et al. 2019



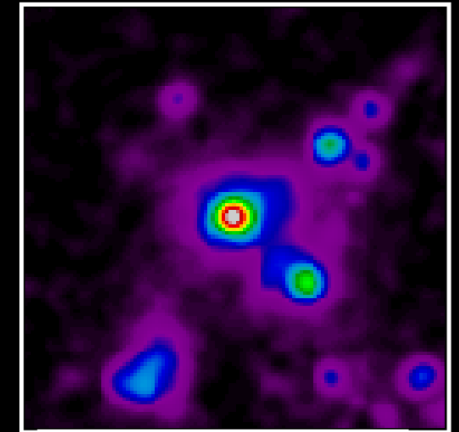
AGN + host  
galaxy



Convolved  
model

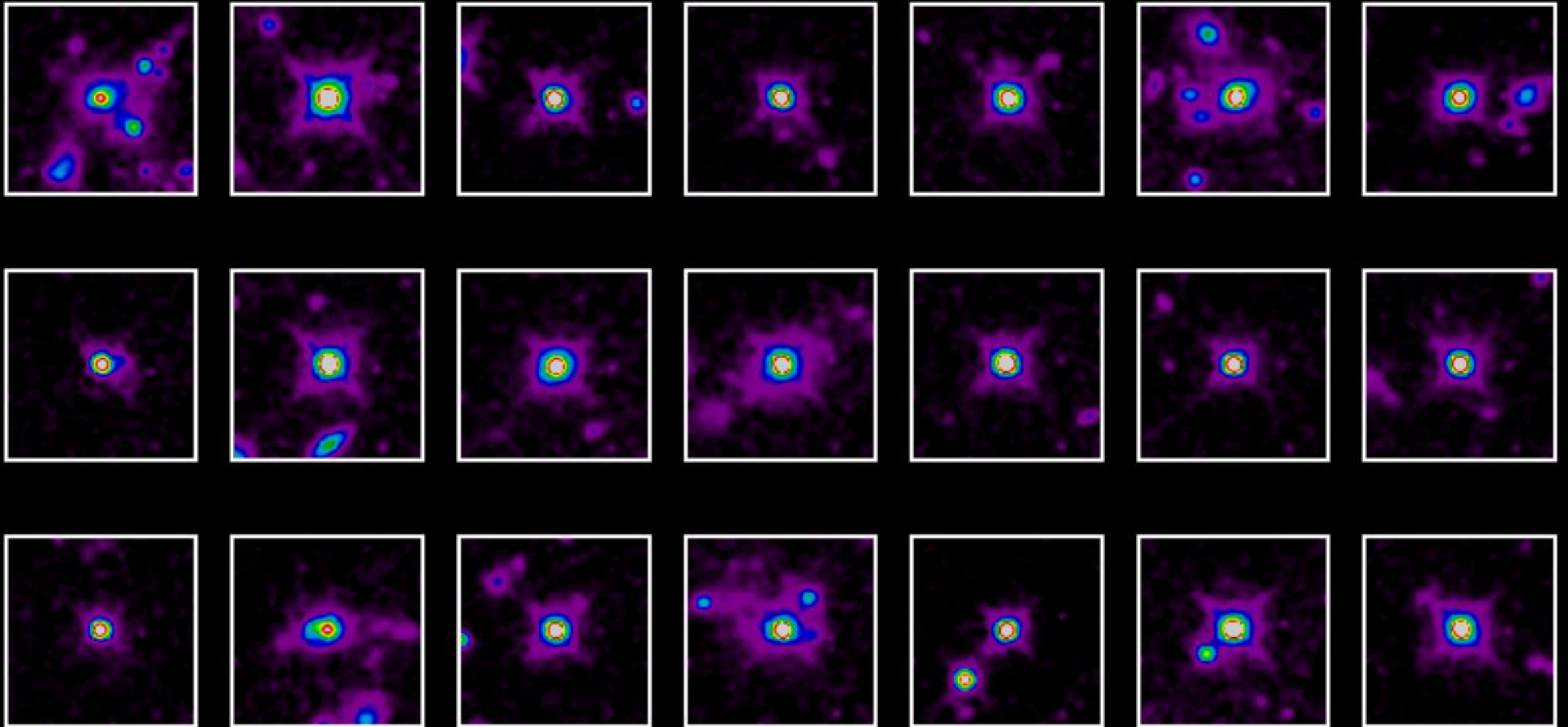


Point-source  
subtracted



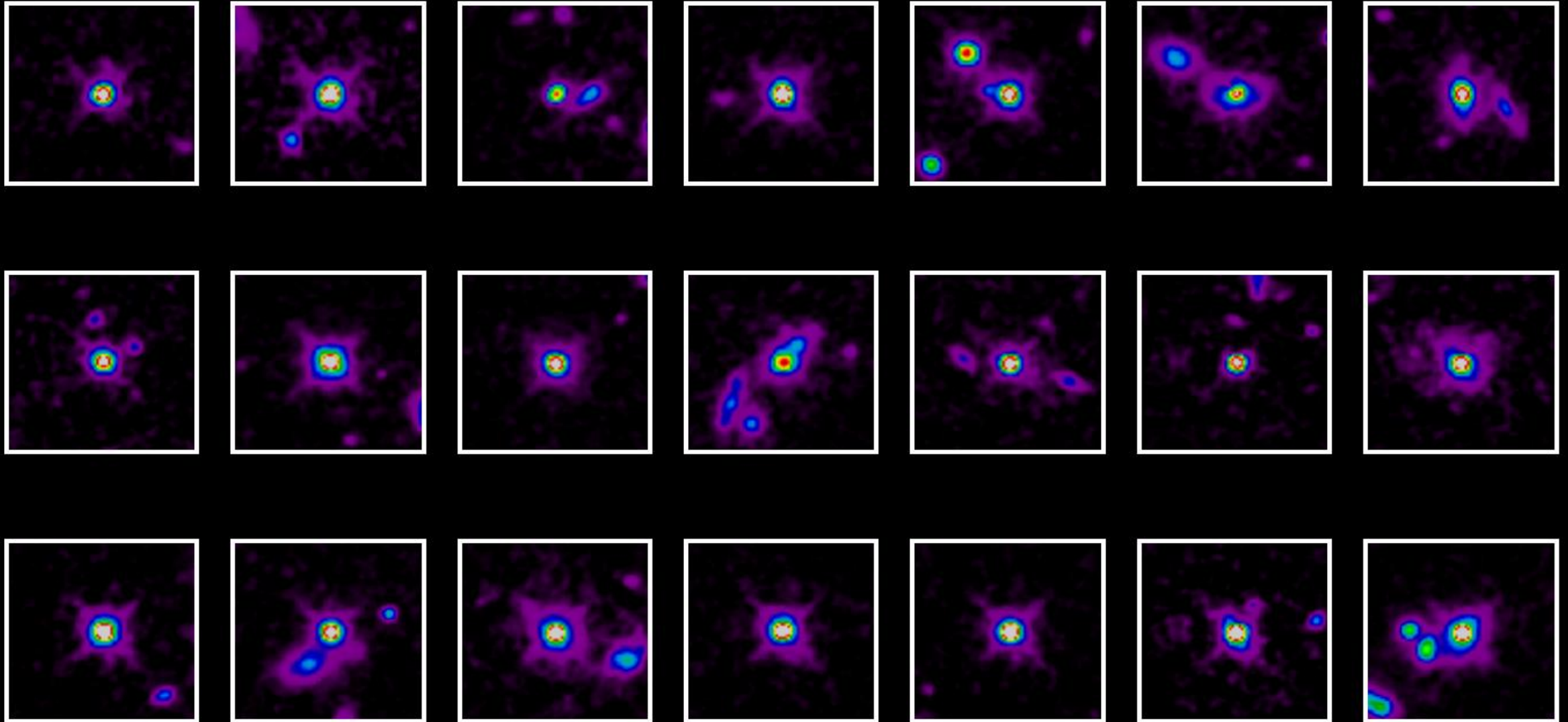
2x2 smoothed

# Modeling & Point-source subtraction



Marian et al. 2019

# Modeling & Point-source subtraction



Marian et al. 2019

# Analysis

Main steps:

- Coauthors rank joint sample
  - ↳ (S. Cohen, B. Husemann, K. Jahnke, V. Jones, A. Koekemoer, V. Marian, A. Schulze, A. van der Wel, C. Villforth & R. Windhorst)
- Combine individual rankings
- Determine cut-off rank
- Split sample
- Derive merger fractions



# Analysis

Main steps:

- Coauthors rank joint sample
  - ↳ (S. Cohen, B. Husemann, K. Jahnke, V. Jones, A. Koekemoer, V. Marian, A. Schulze, A. van der Wel, C. Villforth & R. Windhorst)
- Combine individual rankings
- Determine cut-off rank
- Split sample
- Derive merger fractions

# Analysis

Main steps:

- Coauthors rank joint sample

↳ (S. Cohen, B. Husemann, K. Jahnke, V. Jones, A. Koekemoer, V. Marian, A. Schulze, A. van der Wel, C. Villforth & R. Windhorst)

- Combine individual rankings

- Determine cut-off rank



*We checked the impact of different combination methods and cut-off ranks!*

- Split sample

- Derive merger fractions

# Analysis

Main steps:

- Coauthors rank joint sample

↳ (S. Cohen, B. Husemann, K. Jahnke, V. Jones, A. Koekemoer, V. Marian, A. Schulze, A. van der Wel, C. Villforth & R. Windhorst)

- Combine individual rankings

- Determine cut-off rank



*We checked the impact of different combination methods and cut-off ranks!*

- Split sample

- Derive merger fractions

# Merger fractions

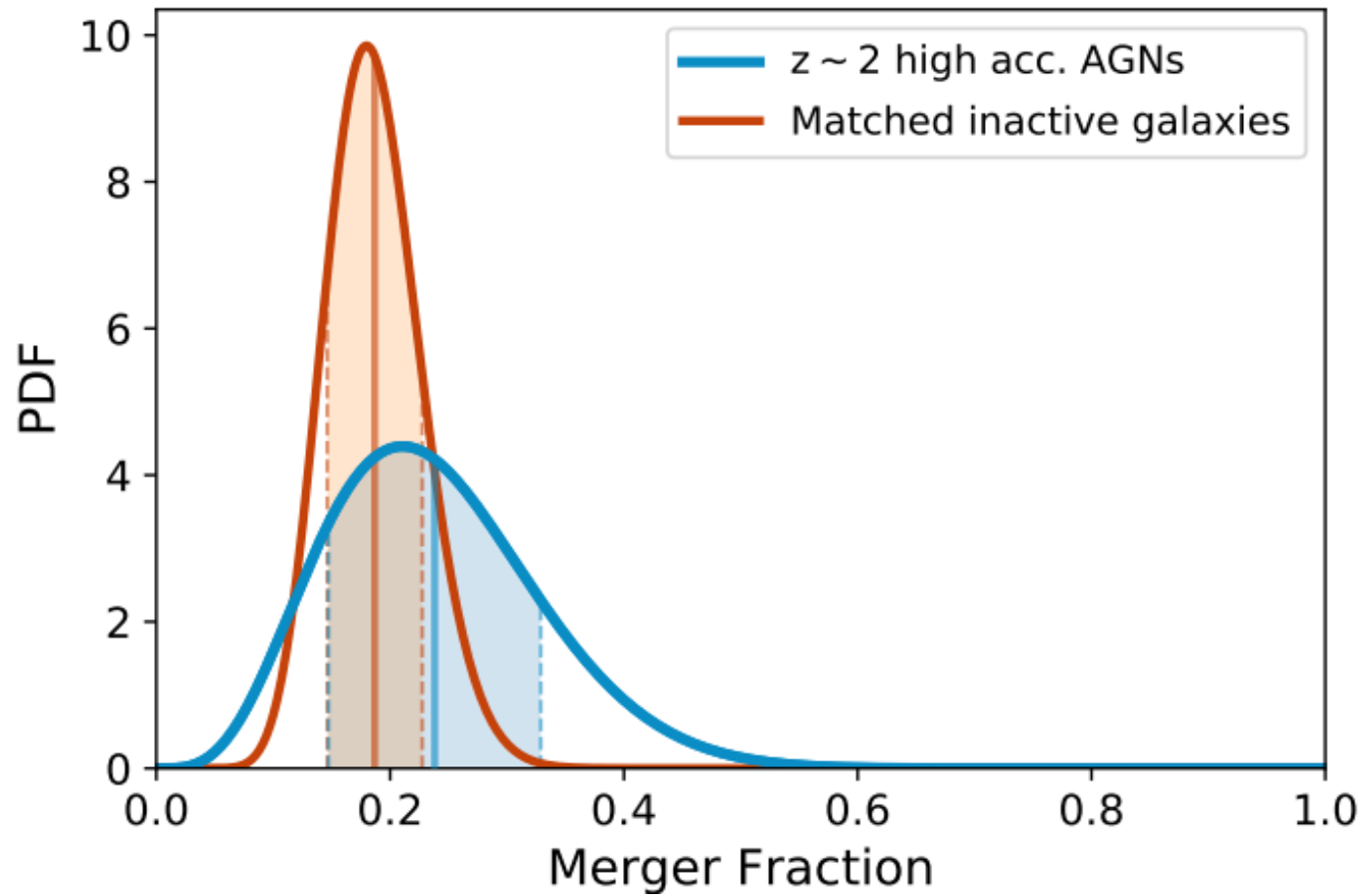
- Showing merger features:
  - 5 AGN host galaxies
  - 17 inactive galaxies
- Showing no such features:
  - 16 AGN host galaxies
  - 74 inactive galaxies

# Merger fractions

- Showing merger features:
  - 5 AGN host galaxies
  - 17 inactive galaxies
- Showing no such features:
  - 16 AGN host galaxies
  - 74 inactive galaxies

$$f_{(m,agn)} = 0.24 \pm 0.09$$

$$f_{(m,ina)} = 0.19 \pm 0.04$$



Marian et al. 2019

$$f_{(m,agn)} = 0.24 \pm 0.09$$

$$f_{(m,ina)} = 0.19 \pm 0.04$$

# But what about...

- ...a dependence on stellar mass?

# But what about...

- ~~...a dependence on stellar mass?~~
- ...a time lag between merger and AGN activity?



# But what about...

- ~~...a dependence on stellar mass?~~
- ...a time lag between merger and AGN activity?
  - Usual lifetimes of AGNs:  $10^6 - 10^8$  yr (Martini 04; Hopkins+ 05; Shen+ 07; Hopkins & Hernquist 09; Conroy & White 13; Cen & Safarzadeh 15)
  - Lifetime of merger features:  $10^9 - 10^{10}$  yr (Conselice 06; Lotz+ 08; Ji+ 14; Solanes+ 18)
  - Visibility overlap of  $\geq 500$  Myr, even with delay of  $\sim 300$  Myr

# But what about...

- ~~...a dependence on stellar mass?~~
- ~~...a time lag between merger and AGN activity?~~
- ... intermittence of AGN activity?

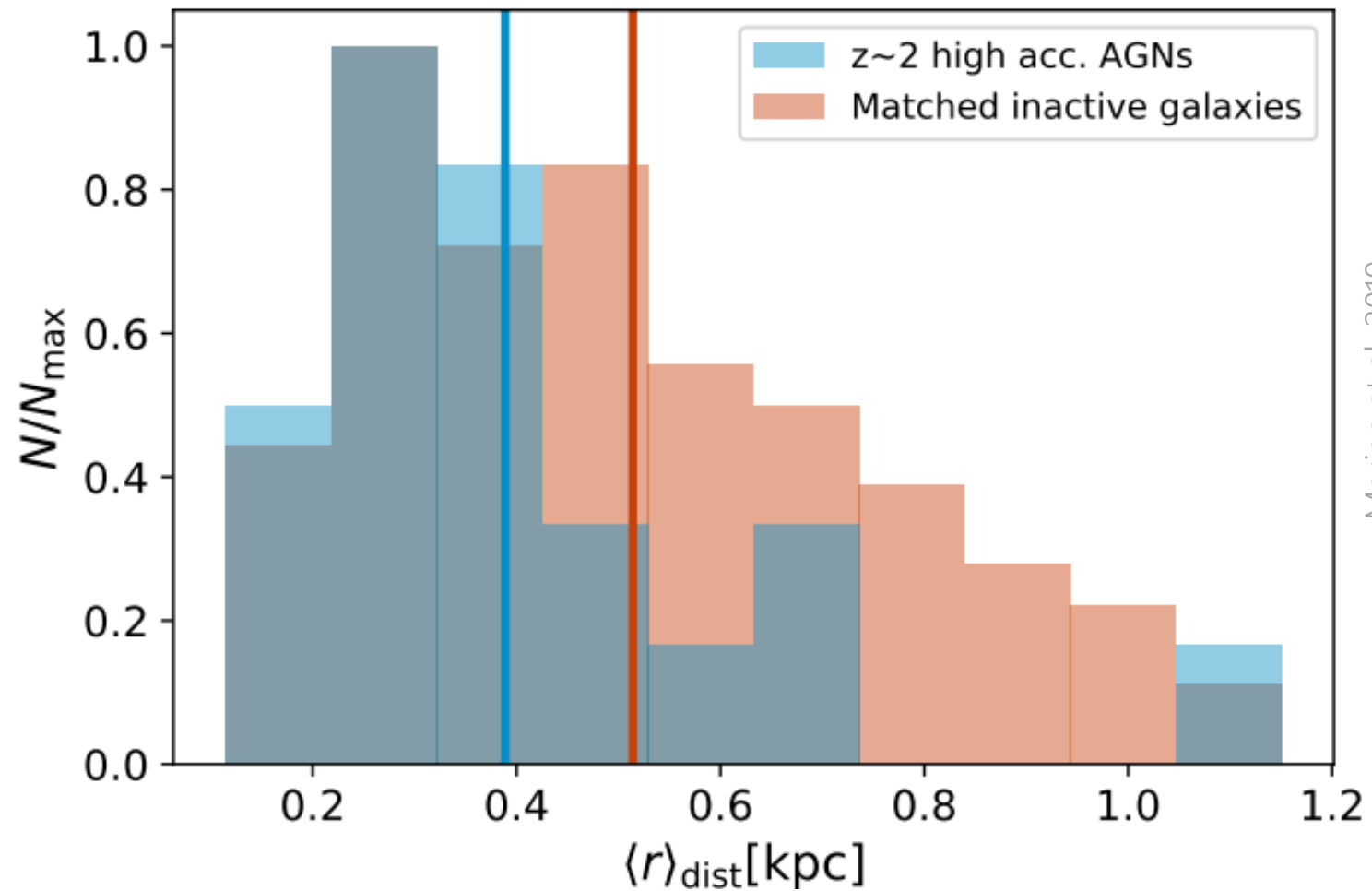
# But what about...

- ~~...a dependence on stellar mass?~~
- ~~...a time lag between merger and AGN activity?~~
- ... intermittence of AGN activity?
  - AGN timescales:  $\sim 10^5$  yr  
or
  - $\sim 20\%$  of merging inactive galaxies host intermittent AGN

# But what about...

- ~~...a dependence on stellar mass?~~
- ~~...a time lag between merger and AGN activity?~~
- ~~... intermittence of AGN activity?~~
- ... a potential offset between AGN position and host galaxy flux center

# Spatial offset between AGN and host galaxy nucleus



Marian et al. 2019

# But what about...

- ~~...a dependence on stellar mass?~~
- ~~...a time lag between merger and AGN activity?~~
- ~~... intermittence of AGN activity?~~
- ~~... a potential offset between AGN position and host galaxy flux center~~

# Take away conclusion

**Major mergers are not the dominant trigger  
for (high-accretion) AGNs (at  $z=2$ )**

# Take away conclusion

**Major mergers are not the dominant trigger  
for (high-accretion) AGNs (at  $z=2$ )**

More details in

Marian V., Jahnke K., Mechtley M., Cohen S., Husemann B., Jones V.,  
Koekemoer A., Schulze A., van der Wel A., Villforth C., Windhorst R.

(subm. to ApJ, arXiv:1904.00037)



# Take away conclusion

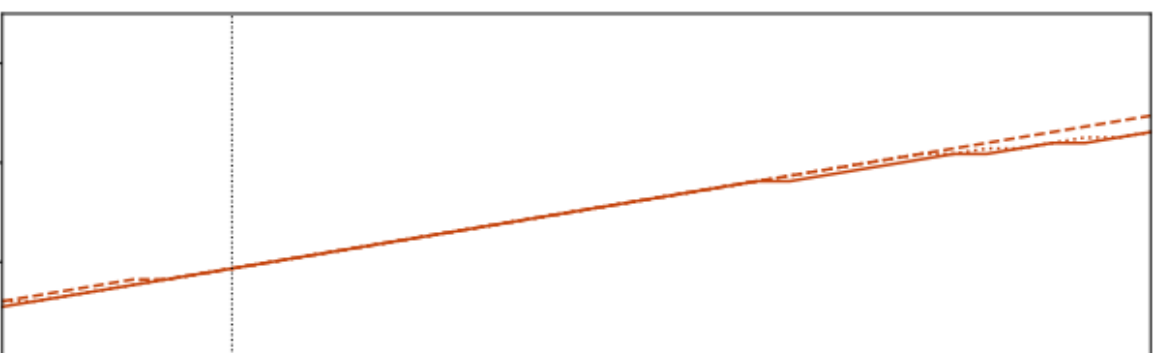
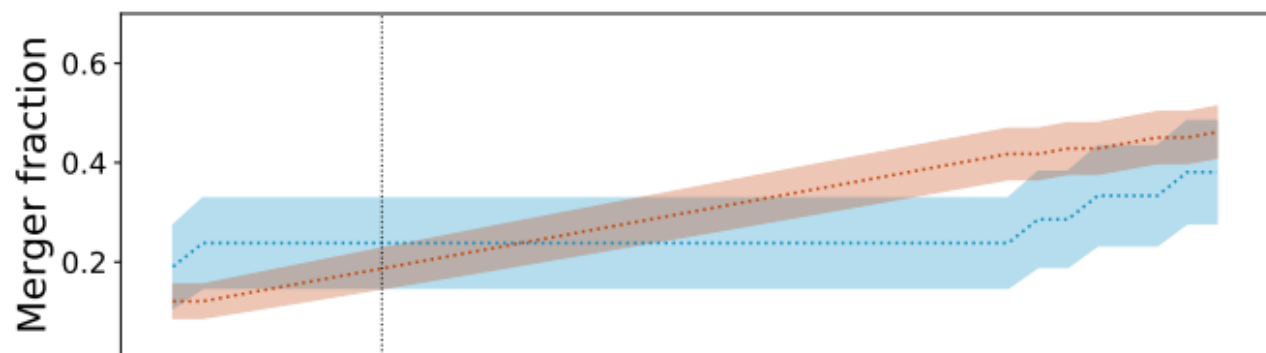
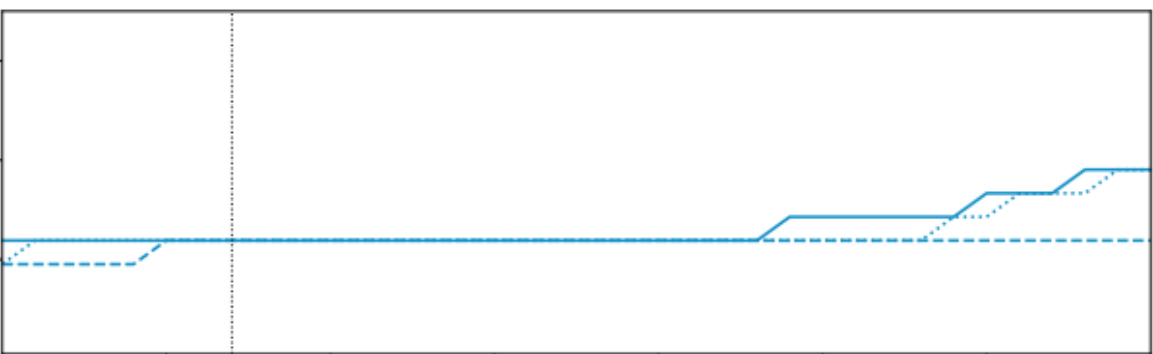
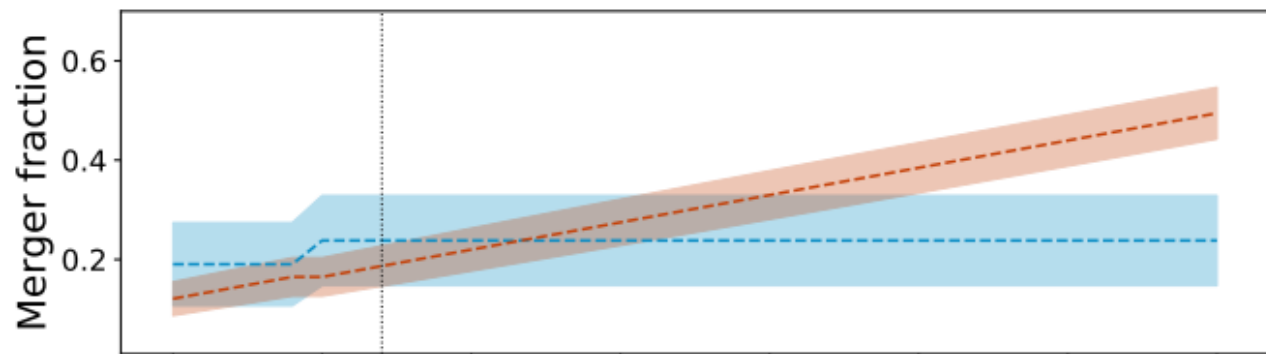
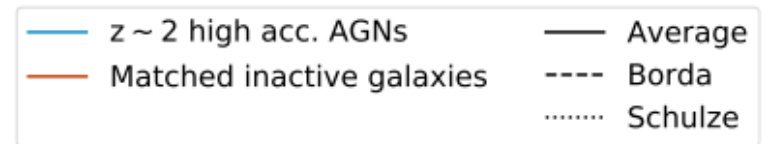
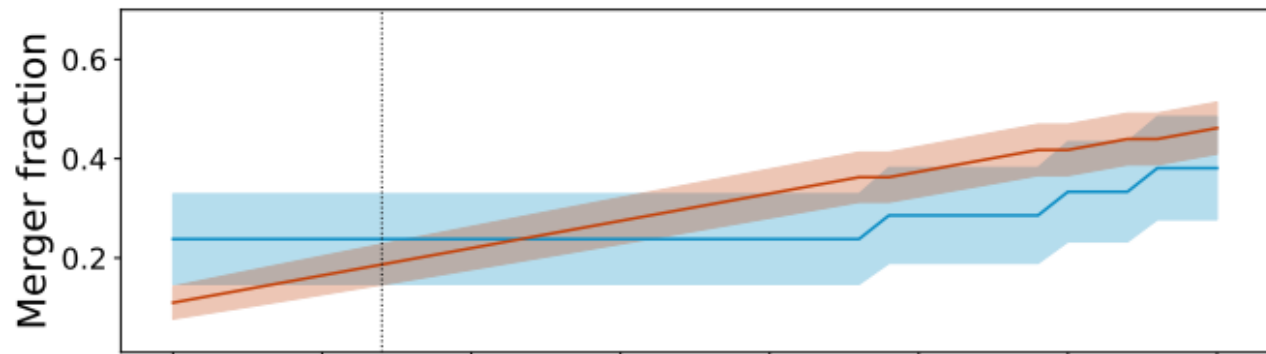
**Major mergers are not the dominant trigger  
for (high-accretion) AGNs (at  $z=2$ )**

More details in

Marian V., Jahnke K., Mechtley M., Cohen S., Husemann B., Jones V.,  
Koekemoer A., Schulze A., van der Wel A., Villforth C., Windhorst R.

(subm. to ApJ, arXiv:1904.00037)

# Thank you



Credit: Marian et al. (2019)

# Intermittent AGN

$$f_{m,ina \& agn} = f_{agn} \times f_{m,agn} \times \frac{t_m}{t_{agn}}$$

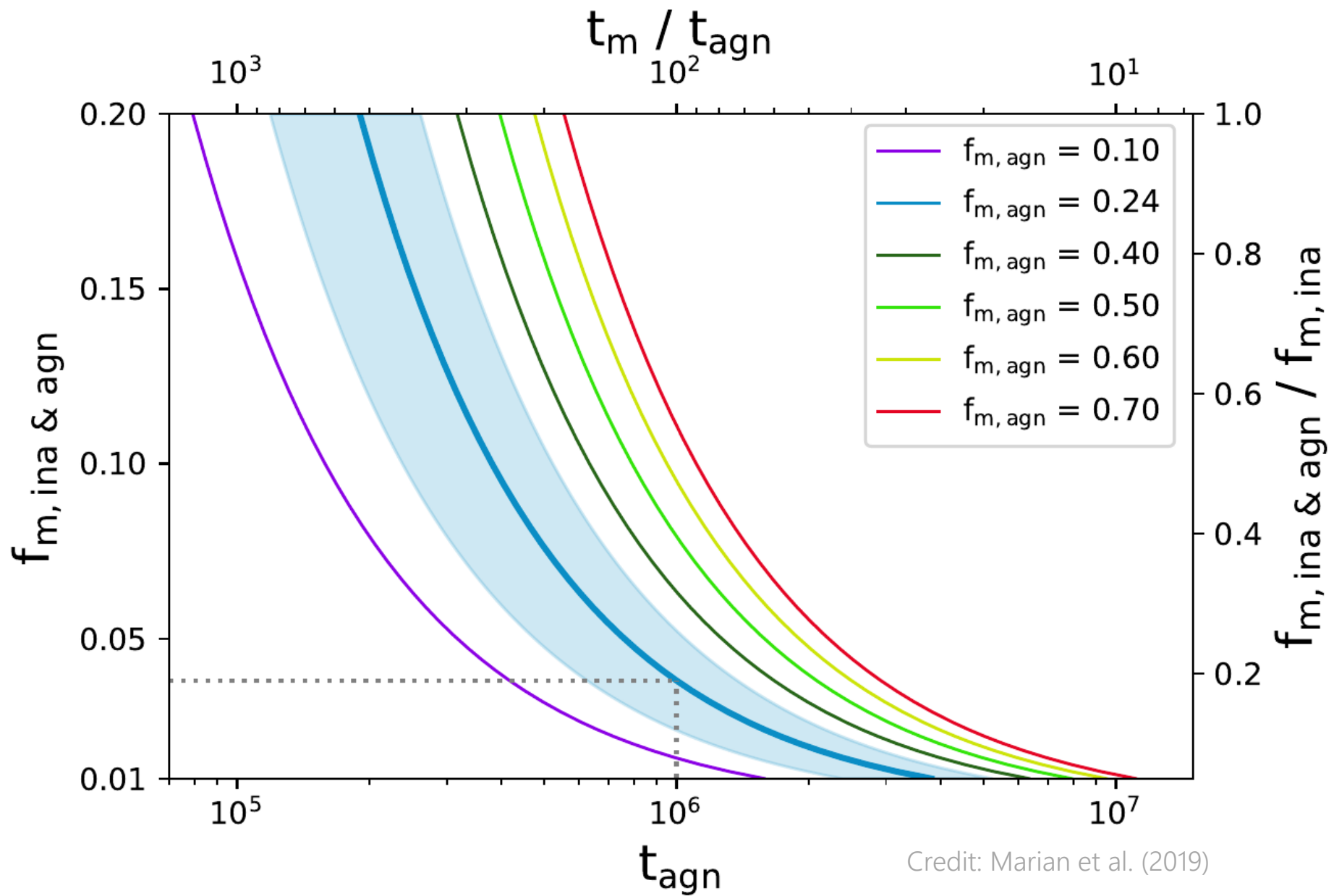
Total fraction of AGN within our constraints

Fraction of merging, inactive galaxies that host an intermittent AGN

AGN merger fraction

Merger timescale

AGN timescale



Credit: Marian et al. (2019)